

# Astrophysics of Reference Frame Tie Objects

Principal Investigator: Kenneth J. Johnston, (USNO)

Team Members:

David Boboltz (USNO), Alan Lee Fey (USNO), Ralph A. Gaume (USNO),  
Norbert Zacharias (USNO).

The Astrophysics of Reference Frame Tie Objects Key Science program will investigate the underlying physics of SIM grid objects. Extragalactic objects in the SIM grid will be used to tie the SIM reference frame to the quasi-inertial reference frame defined by extragalactic objects and to remove any residual frame rotation with respect to the extragalactic frame. The current realization of the extragalactic frame is the International Celestial Reference Frame (ICRF). The ICRF is defined by the radio positions of 212 extragalactic objects and is the IAU sanctioned fundamental astronomical reference frame. This key project will advance our knowledge of the physics of the objects which will make up the SIM grid, such as quasars and chromospherically active stars, and relates directly to the stability of the SIM reference frame. The following questions concerning the physics of reference frame tie objects will be investigated.

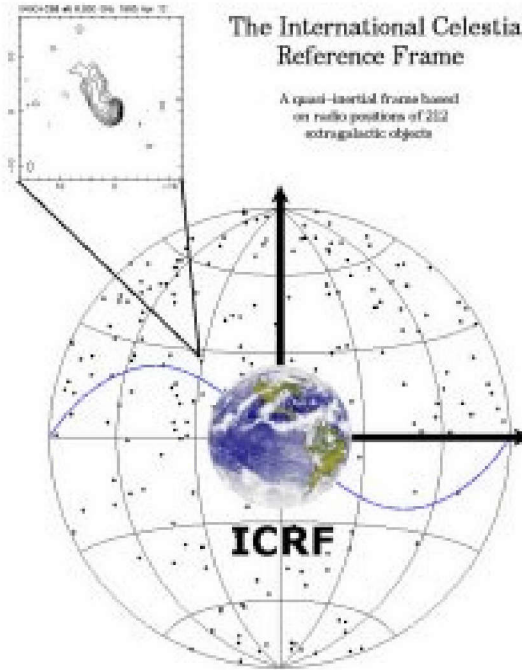


Figure 1: Figure shows a schematic representation of the ICRF as a projection of extragalactic radio sources onto the celestial sphere. The inset shows the 8.55 GHz radio frequency image of a representative source.

What is the origin of optical emission in quasars? In the standard theory of extragalactic radio sources, emission from quasars and active galactic nuclei (AGN) is assumed to be powered by a central engine (presumably a black hole) where energetic phenomena occur. The origin of the optical wavelength radiation from AGN identified with compact radio sources is not well

Report Documentation Page				Form Approved OMB No. 0704-0188	
Public reporting burden for the collection of information is estimated to average 1 hour per response, including the time for reviewing instructions, searching existing data sources, gathering and maintaining the data needed, and completing and reviewing the collection of information. Send comments regarding this burden estimate or any other aspect of this collection of information, including suggestions for reducing this burden, to Washington Headquarters Services, Directorate for Information Operations and Reports, 1215 Jefferson Davis Highway, Suite 1204, Arlington VA 22202-4302. Respondents should be aware that notwithstanding any other provision of law, no person shall be subject to a penalty for failing to comply with a collection of information if it does not display a currently valid OMB control number.					
1. REPORT DATE <b>2005</b>		2. REPORT TYPE <b>N/A</b>		3. DATES COVERED <b>-</b>	
4. TITLE AND SUBTITLE <b>Astrophysics of Reference Frame Tie Objects</b>				5a. CONTRACT NUMBER	
				5b. GRANT NUMBER	
				5c. PROGRAM ELEMENT NUMBER	
6. AUTHOR(S)				5d. PROJECT NUMBER	
				5e. TASK NUMBER	
				5f. WORK UNIT NUMBER	
7. PERFORMING ORGANIZATION NAME(S) AND ADDRESS(ES) <b>Library U.S. Naval Observatory 3450 Massachusetts Avenue, N.W. Washington, D.C. 20392-5420</b>				8. PERFORMING ORGANIZATION REPORT NUMBER	
9. SPONSORING/MONITORING AGENCY NAME(S) AND ADDRESS(ES)				10. SPONSOR/MONITOR'S ACRONYM(S)	
				11. SPONSOR/MONITOR'S REPORT NUMBER(S)	
12. DISTRIBUTION/AVAILABILITY STATEMENT <b>Approved for public release, distribution unlimited</b>					
13. SUPPLEMENTARY NOTES					
14. ABSTRACT					
15. SUBJECT TERMS					
16. SECURITY CLASSIFICATION OF:			17. LIMITATION OF ABSTRACT <b>UU</b>	18. NUMBER OF PAGES <b>3</b>	19a. NAME OF RESPONSIBLE PERSON
a. REPORT <b>unclassified</b>	b. ABSTRACT <b>unclassified</b>	c. THIS PAGE <b>unclassified</b>			

established. There are several possibilities: 1) the optical radiation is thermal emission from an accretion disk; 2) the optical radiation is non-thermal emission from a magnetized corona or wind emanating from the central region of the accretion disk; or 3) the optical radiation is emission from a relativistic jet beamed toward the observer. SIM observations will allow a direct test of this model.

Are the optical photo-centers of quasars compact and positionally stable on the micro-arcsecond level? It is well known that variable intrinsic structure of AGN at radio wavelengths has a significant impact on (degrades) radio positional accuracy. Currently, there is very little available information on whether the optical counterparts of the radio objects are compact at the level of astrometric precision expected from SIM. Although the core of a quasar's optical emission may originate in a region as small as 1 pc (200 microarcseconds at 1 Gpc), some degree of photocenter wander should be expected, probably correlated with optical variability. Motion of a quasar's photocenter may also result from a variable nucleus in combination with effects in the larger (albeit fainter) host galaxy.

Are binary black hole mergers responsible for quasars? Optical photocenter wander of quasars can be used to search for binary black hole signatures and thus test the hypothesis that binary black hole mergers are responsible for most quasars. A positive detection of binary black hole induced motions in the optical photocenter of quasars would be good evidence to support this hypothesis.

What is(are) the emission mechanism(s) responsible for generating radio emission in chromospherically active stars. Is the emission thermal, relativistic synchrotron or gyro-synchrotron? No consensus has yet developed concerning the physics of the formation and evolution of the radio emission associated with the active binary star systems. Also, for most active binaries, the location of the radio emission with respect to the binary components is unknown; e.g. is the radio emitting region centered on one of the stars, is it located in the intra-binary region, or does it surround both stars? This uncertainty can be attributed, in part, to inadequacies in the radio/optical frame link.

What causes the transition of spherically symmetric Asymptotic Giant Branch (AGB) stars to asymmetric planetary nebulae (PNe)? Models based on interacting winds have had considerable success explaining PNe morphologies. However, it is still unknown whether the asymmetries observed in PNe are present in the progenitor AGB stars themselves, and whether these asymmetries could be produced by mechanisms such as non-radial pulsations or unseen companions. SIM observations of a number of AGB stars shown to have asymmetric envelopes, through radio observations of circumstellar masers, will allow a test of these theories.

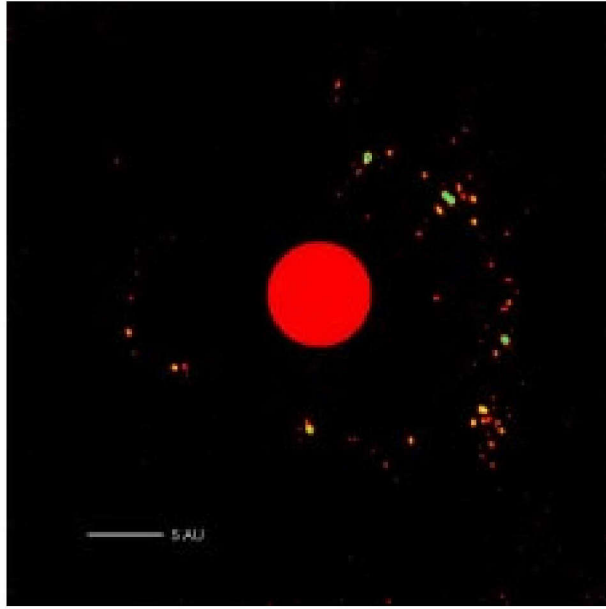


Figure 2: Figure shows the asymmetrical SiO maser shell toward the Mira variable IK Tauri with nearly a 2:1 axial ratio. The location of the stellar disk (red circle) relative to the masers is pure conjecture.